



SiPM development for Astroparticle Physics applications

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Abstract: The SiPM is a novel solid state photodetector which can be operated in the single photon counting mode. It has excellent features, such as high quantum efficiency, good charge resolution, fast response, very compact size, high gain of 10^6 , very low power consumption, immunity to the magnetic field and low bias voltage (30-70V). Drawbacks of this device currently are a large dark current, crosstalk between micropixels and relatively low sensitivity to UV and blue light. In the last few years, we have developed large size SiPMs (9 mm² and 25 mm²) for applications in the imaging atmospheric Cherenkov telescopes, MAGIC and CTA, and in the space-borne fluorescence telescope EUSO. The current status of the SiPM development by MPI and MPEHI will be presented.

Introduction

In Astroparticle Physics experiments, we measure the rare and low flux events, and naturally the instrument will become large in area and volume. Photodetection techniques play an important role in the detection of these rare and low flux events. We require low level light (LLL) detectors working in single photon counting mode. We record the time profile of signals or require the very fast timing of signals. The imaging of the signal plays a very important role in Imaging Air Cherenkov Telescopes (IACTs) for high energy gamma rays and in the air fluorescence technique for Ultra High Energy Cosmic Rays (UHECRs).

The SiPM is a solid state photodetector developed for applications in High Energy Physics, Astroparticle Physics and Medical Science. For Astroparticle Physics applications, the most attractive feature of the SiPM is its high photodetection efficiency ($\sim 70\%$) and the capability of single photon counting. Further, there are many other advantages, such as the compactness of its size and volume, the low bias voltages, the very high gain of 10^6 , the low power consumption, etc.. At present, the major problems of SiPMs are the small size, the cross talk between micropixels in the SiPM and the high dark current.

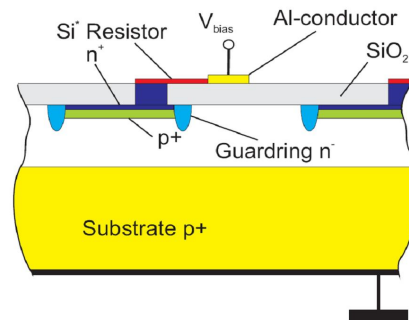


Figure 1: Structure of the micropixel of a SiPM.

For the MAGIC telescope [1], the high photodetection efficiency will allow us to lower the threshold energy of gamma ray detection down to 10-20 GeV. We can also design new smaller and inexpensive Cherenkov telescopes keeping the performance at an equally high level. This allows us to consider an array of smaller high performance telescopes at a low budget. For MAGIC, the pixel size of the photodetector shall be a 10 – 20 mm size (corresponding to $0.05 - 0.1^\circ$ with a winston cone). Even if we consider smaller pixels below 0.05° , we do not gain in angular resolution and gamma/hadron separation.

For the space-borne air fluorescence detector EUSO[2], photodetectors with a size of 3 – 5 mm

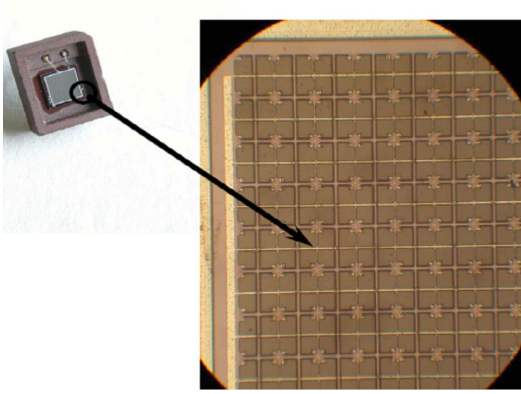


Figure 2: Microscope picture of a $9(3 \times 3)$ mm² SiPM. It consists of an array of 5625(75×75) micropixels with 30×30 μm^2 size.

are required for 0.1° pixel imaging. The high photodetection efficiency will ensure the detection of UHECRs above $(2-3) \times 10^{19}$ eV and allow a large overlap in the energy range with ground based detectors, Auger[4] and TA[5]. Alternatively it will allow us to use the high altitude orbit or tilted mode to cover larger detection volumes for UHECRs. Characteristics like the low power consumption and insensitivity to magnetic fields are important advantages for space detectors.

Therefore, the larger size SiPM of 25 mm² or even 100 mm², with high photodetection efficiency, low dark rate, low crosstalk and enhanced sensitivity in UV and blue light is the prime target of our SiPM development. Silicon itself has the nice feature of nearly 100% internal quantum efficiency. The increase of the filling factor of the sensitive area on the device surface and the increase of the Geiger efficiency will give us a very high photodetection efficiency.

Structure and operation of the SiPM

The SiPM consists of an array of a large number of micropixels (avalanche diodes) working in Geiger mode. The schematic picture of the individual micropixel's structure is shown in figure 1. This micropixel works somehow in digital mode, giving us the signal $Q = C_{pixel} \times (V - V_{bd})$ in each Geiger discharge. Here, C_{pixel} is the capacitance between anode and cathode of the avalanche diode, V_{bd} is

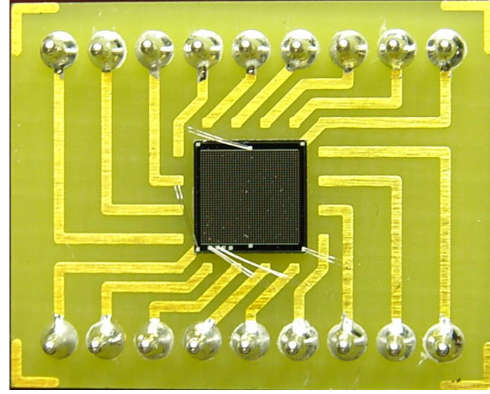


Figure 3: 25(5×5) mm² SiPM. It consists of an array of 1600(40×40) micropixels with 100×100 μm^2 size.

the breakdown voltage of micropixel (avalanche diode), $V - V_{bd}$ is called over voltage. The typical Q is $10^6 e$. The anodes of all micropixels are joined together on a common substrate and are working on common load in case of n on p structure.

The operation of the individual micropixel is as follows. The micropixel is biased with a higher voltage than the breakdown voltage. This over voltage is typically 10% or 20% of the breakdown voltage. The micropixel is then kept in a meta-stable condition, and if a pair of electron and hole is created in the depletion or drift region by an incident photon, a Geiger discharge will be triggered. The current flow from the biasing circuit through the quenching register causes the bias voltage to drop and the Geiger process will stop automatically. The biasing circuit will recharge the capacitance of the SiPM to raise the bias voltage again up to the meta-stable over voltage. This recharge takes typically 1 μsec . The dark rate (night sky back ground, or dark current) DR per micropixel of 30×30 μm^2 or 100×100 μm^2 is usually low enough, i.e. 1 – 10 kHz, even at room temperature, so the dead time and the decrease of photodetection efficiency can be negligibly small.

Development of SiPMs for MAGIC, CTA and EUSO

MPI and MEPhI have developed SiPMs with a size of $9(3 \times 3)$ mm² and $25(5 \times 5)$ mm² for MAGIC [1], CTA [3] and EUSO [2] applications as shown

in figure 2 and figure 3[6,7,8,9]. The micropixel size is enlarged to $100 \times 100 \mu\text{m}^2$ to get a higher filling factor of the sensitive region, in case of the 25 mm^2 SiPM. The devices were manufactured by the company Pulsar in Moscow.

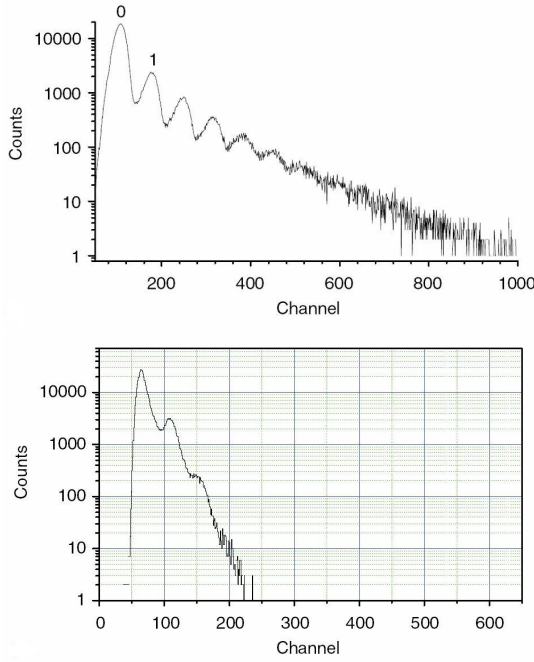


Figure 4: Top and bottom panels show the pulse height distributions of dark current without and with OC suppression structure, respectively. They are operated with the gain of 3×10^6 and 3×10^7 . In both panels, the first peak is pedestal, the second and third one are single and two ph.e., respectively.

Dark Rate

The dark rate of SiPMs is about $1 - 2 \text{ MHz/mm}^2$ at room temperature ($\sim 20^\circ\text{C}$). Dark current is produced by two mechanisms, i.e. the electric field assisted tunneling effect and the thermal effect. The dark current can be reduced by one or two orders of magnitude by cooling the temperature down to -7°C or -35°C in our devices, respectively. In our device the main component of dark current is the thermal one. The clever design avoids unnecessary extremely high electric fields will reduce the electric field assisted tunneling. The usage of ultra pure silicon wafers, the reduction of the volume of

depletion and drift regions will help to reduce the dark rate.

Optical Crosstalk

The Optical Crosstalk (OC) is a specific phenomenon in SiPMs. The Geiger avalanche in one micropixel will provoke the emission of light, typically ~ 10 photons/avalanche of $10^6 e$, and these photons will hit other pixels. This means that the single ph.e. has a chance to produce a multiple ph.e. signal. The emission of light is proportional to the Geiger gain. The Geiger efficiency which determines the final photodetection efficiency has also a strong positive dependence on the Geiger gain. Therefore, if we want to obtain a high photodetection efficiency, we need to operate SiPMs in high Geiger gain. This condition will introduce severe optical crosstalk in SiPMs. To avoid this problem we have introduced the OC suppression structure.

Optical crosstalk entails a serious problem in our specific usage, because we use photodetectors not only to record photon intensity but also to trigger events. Especially in the Cherenkov experiment and air fluorescence experiment, we have a high rate of background photons from the night sky. Even if there are a series of single ph.e., this optical crosstalk produces multiple ph.e. signals with a high rate and we are consequently required to effectively set the trigger threshold high in order to obtain a reasonable trigger rate.

The top panel of figure 4 shows the pulse height distribution of dark rate in the standard SiPM without the OC suppression structure. We can see the exponential long tail in the multiple ph.e. region. These large pulses will raise the trigger threshold like after pulses in PMTs. By manufacturing the OC suppression structure (the groove structures between micropixels which block the transmission of light between micropixels through silicon), we succeed to suppress these unfavorable optical crosstalk. As shown in the bottom of figure 4, the optical crosstalk level was significantly reduced by this improvement. We can evaluate this improvement with the parameter of the Excess Noise Factor (ENF); the ENF of about 1.6 was improved to $\text{ENF} = 0.97 \pm 0.05$ with the OC suppression structure.

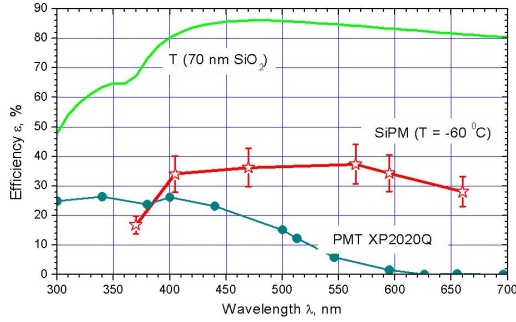


Figure 5: The photodetection efficiency of SiPM of $25(5 \times 5) \mu\text{m}^2$ with micropixels size of $100 \times 100 \mu\text{m}^2$ (filling factor of 64%) at a temperature of -60°C is compared with that of a Photomultiplier.

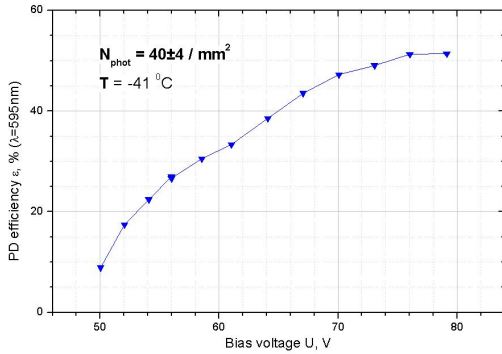


Figure 6: The photodetection efficiency of SiPM of $25(5 \times 5) \mu\text{m}^2$ with micropixels size of $100 \times 100 \mu\text{m}^2$ (filling factor of 64%) at large overvoltages.

Photodetection efficiency

Photodetection efficiency PDE is determined by four factors, i.e. the light transmission from surface to depletion or drift layers T , the filling factor of the sensitive area, (sensitive area)/(all area) F , the quantum efficiency of silicon $QE(\lambda)$, and the Geiger efficiency G .

$$\text{PDE} = T \times F \times QE(\lambda) \times G.$$

Silicon has a high reflectivity index $n \sim 4.0$, and with AR coating we can increase the light transmission up to a 90–95% level. As shown in figure 1, there are guard rings, quenching registers, aluminum conductors for bias voltage supply around micropixels. This structure will limit the filling

factor to 80%. The internal quantum efficiency of silicon $QE(\lambda)$ is close to 100%, and the Geiger efficiency can be raised up to $\sim 90\%$ in high gain operation (high over voltage). Theoretically, the total PDE can be increased up to $\sim 70\%$ with the existing technologies.

The PDE of a 9 mm^2 SiPM is about 25% between 450nm and 550nm[6]. This 9 mm^2 SiPM consists of 5625 ($= 75 \times 75$) micropixels with a size of $30 \times 30 \mu\text{m}^2$ (filling factor of $\sim 50\%$).

The SiPM of 25 mm^2 consists of 1600 (40×40) micropixels with the size of $100 \times 100 \mu\text{m}^2$ (the filling factor of 64%). The PDE of such a structure is shown in figure 5. For large overvoltages where the Geiger efficiency is close to 100%, the PDE of a $100 \times 100 \mu\text{m}^2$ structure reaches $\sim 50\%$ as shown in figure 6.

Conclusion

We have successfully developed 9 mm^2 and 25 mm^2 SiPMs. Their characteristics are almost satisfactory, but we consider further developments/improvements, the enhancement of UV-Blue light sensitivity with reverse structure (p on n), and larger SiPMs of 100 mm^2 size or the array of 4×4 SiPMs of 25 mm^2 size for MAGIC [1], CTA [3] and EUSO [2] applications.

Acknowledgements

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Reference

- [1] <http://www.magic.mppmu.mpg.de/>
- [2] <http://euso.riken.go.jp/>
- [3] see contributions in this conference
- [4] <http://www.auger.org/>
- [5] <http://www-ta.icrr.u-tokyo.ac.jp/>
- [6] B. Dolgoshein et al., NIM A563 (2006) 368.
- [7] R. Mirzoyan et al., N.I.M. A572 (2007) 493.
- [8] P. Buzhan et al., N.I.M. A567 (2006) 78.
- [9] A.N. Otte et al., N.I.M. A 545 (2005) 705.